100 years of black hole physics

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Outline

1. A century-old history
2. Black holes in the sky
3. Observing black holes via gravitational waves: a dream come true
4. Testing general relativity with black holes
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3. Observing black holes via gravitational waves: a dream come true

4. Testing general relativity with black holes
A century-old history

A two centuries-old prehistory...

\[ V_{\text{esc}} > c \iff \frac{2GM}{R} > c^2 \iff \frac{2G}{R} \times \frac{4}{3} \pi R^3 \rho > c^2 \iff R > \sqrt{\frac{3c^2}{8\pi G \rho}} \]
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John Michell (1784)

“If there should really exist in nature any bodies, whose density is not less than that of the sun, and whose diameters are more than 500 times the diameter of the sun, since their light could not arrive at us, ..., we could have no information from sight”

[Phil. Trans. R. Soc. Lond. 74, 35 (1784)]
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Pierre Simon de Laplace (1796)

“Un astre lumineux, de la même densité que la Terre, et dont le diamètre serait 250 fois plus grand que le Soleil, ne permettrait, en vertu de son attraction, à aucun de ses rayons de parvenir jusqu’à nous. Il est dès lors possible que les plus grands corps lumineux de l’univers puissent, par cette cause, être invisibles.”

[Exposition du système du monde (1796)]
Limits of the Newtonian concept of a black hole

- No privileged role of the velocity of light in Newtonian theory: nothing forbids $V > c$: the “dark stars” are not causally disconnected from the rest of the Universe.
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- $V_{\text{esc}} \sim c \implies$ gravitational potential energy $\sim$ mass energy $Mc^2$.

$\implies$ a relativistic theory of gravitation is necessary!
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\[ \implies \text{a relativistic theory of gravitation is necessary!} \]

- No clear action of the gravitation field on electromagnetic waves in Newtonian gravity

[R. Taillet]
100 years ago: a relativistic theory of gravitation

844 Sitzung der physikalisch-mathematischen Klasse vom 25. November 1915

Die Feldgleichungen der Gravitation.
Von A. Einstein.

In zwei vor kurzem erschienenen Mitteilungen habe ich gezeigt, wie man zu Feldgleichungen der Gravitation gelangen kann, die dem Postulat allgemeiner Relativität entsprechen, d. h. die in ihrer allgemeinen Fassung beliebigen Substitutionen der Raumzeitvariablen gegenüber kovariant sind.

\[ R - \frac{1}{2} R g = \frac{8\pi G}{c^4} T \]

The Schwarzschild solution

- **Nov-Dec. 1915**: Karl Schwarzschild: first exact non-trivial solution of Einstein equation \( \Rightarrow \) spacetime metric outside a **spherical body** of mass \( M \)

\[
g_{\alpha\beta}dx^\alpha dx^\beta = - \left(1 - \frac{2GM}{c^2r}\right) c^2 dt^2 + \left(1 - \frac{2GM}{c^2r}\right)^{-1} dr^2 + r^2 (d\theta^2 + \sin^2\theta d\phi^2)
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- **1916**: Johannes Droste: circular orbit of photons at \( r = 3GM/c^2 \)
- **1920**: Alexander Anderson: light cannot emerge from the region \( r < R_S := \frac{2GM}{c^2} \) ("shrouded in darkness")
A century-old history

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- 1920: Alexander Anderson: light cannot emerge from the region \( r < R_S := \frac{2GM}{c^2} \) ("shrouded in darkness")
- 1923: George Birkhoff: outside any \textit{spherical} body, the metric is Schwarzschild metric
1932: Georges Lemaître: the singularity at $r = R_S$ is a coordinate singularity: the metric components are regular in Lemaître coordinates $(\tau, \chi, \theta, \varphi)$:

$$g_{\alpha\beta} dx^\alpha dx^\beta = -c^2 d\tau^2 + \frac{R_S}{r} d\chi^2 + r^2 (d\theta^2 + \sin^2 \theta d\varphi^2)$$

$$r = r(\tau, \chi) := \left[ \frac{3}{2} \sqrt{R_S} (c\tau - \chi) \right]^{2/3}$$
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1939: Robert Oppenheimer & Hartland Snyder: gravitational collapse of a homogeneous dust cloud (special case of Lemaître’s general solution)

$\implies$ for an external observer, $R \to R_S$ as $t \to +\infty$
1950: John L. Synge, 1960: Martin Kruskal, George Szekeres: complete mathematical description of Schwarzschild spacetime ($\mathbb{R}^2 \times S^2$ manifold)

Schwarzschild-Droste coordinates $(t, r)$
A century-old history

The Schwarzschild spacetime: Carter-Penrose diagram

figure: http://sagemanifolds.obspm.fr

Éric Gourgoulhon (LUTH)
Rotation enters the game: the Kerr solution

Roy Kerr (1963)

\[
\begin{align*}
g_{\alpha\beta} \, d\alpha \wedge d\beta &= -\left(1 - \frac{2GMr}{c^2\rho^2}\right) \, c^2 dt^2 - \frac{4GMar \sin^2 \theta}{c^2 \rho^2} \, c \, dt \, d\varphi + \frac{\rho^2}{\Delta} \, dr^2 \\
&\quad + \rho^2 d\theta^2 + \left(r^2 + a^2 + \frac{2Ga^2r \sin^2 \theta}{c^2 \rho^2}\right) \sin^2 \theta \, d\varphi^2
\end{align*}
\]

where

\[
\begin{align*}
\rho^2 &:= r^2 + a^2 \cos^2 \theta, \quad \Delta := r^2 - \frac{2GM}{c^2} r + a^2, \quad a := \frac{J}{cM}
\end{align*}
\]

→ 2 parameters: \( M \): gravitational mass; \( J \): angular momentum
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Schwarzschild as the subcase \( a = 0 \):

\[ g_{\alpha\beta} \, dx^\alpha \, dx^\beta = - \left(1 - \frac{2GM}{c^2 r}\right) c^2 dt^2 + \left(1 - \frac{2GM}{c^2 r}\right)^{-1} \, dr^2 + r^2 \left(d\theta^2 + \sin^2 \theta \, d\varphi^2\right) \]
Physical meaning of the parameters $M$ and $J$

- **mass $M$**: not a measure of the “amount of matter” inside the black hole, but rather a *characteristic of the external gravitational field* → measurable from the orbital period of a test particle in far circular orbit around the black hole (*Kepler’s third law*)
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*Remark*: the radius of a black hole is not a well defined concept: it *does not* correspond to some distance between the black hole “centre” and the event horizon. A well defined quantity is the area of the event horizon, $A$. The radius can be then defined from it: for a Schwarzschild black hole:

$$R := \sqrt[4]{\frac{A}{4\pi}} = \frac{2GM}{c^2} \simeq 3 \left( \frac{M}{M_\odot} \right) \text{ km}$$
Slice $t = \text{const}$ and $\theta = \pi/2$ of the Kerr spacetime

Kerr spacetime

- ergosphere
- event horizon
- Cauchy horizon
- singularity
- rotation axis
- equatorial plane

Slice $t = \text{const}$ and $\theta = \pi/2$ of the Kerr spacetime
1964 : Edwin Salpeter, Yakov Zeldovich : quasars (just discovered !) shine thanks to accretion onto a supermassive black hole

1965 : Roger Penrose : if a trapped surface is formed in a gravitational collapse and matter obeys some energy condition, then a singularity will appear

1967 : John A. Wheeler coined the word *black hole*

1969 : Roger Penrose : energy can be extracted from a rotating black hole

1972 : Stephen Hawking : law of area increase $\Rightarrow$ BH thermodynamics

1975 : Stephen Hawking : *Hawking radiation*

1965-1972 : the *no-hair theorem*
Within 4-dimensional general relativity, a stationary black hole in an otherwise empty universe is necessarily a Kerr-Newmann black hole, which is an electro-vacuum solution of Einstein equation described by only 3 parameters:

- the total mass $M$
- the total specific angular momentum $a = J/(Mc)$
- the total electric charge $Q$

$\Rightarrow \text{“a black hole has no hair” (John A. Wheeler)}$
Dorochkevitch, Novikov & Zeldovitch (1965), Israel (1967), Carter (1971), Hawking (1972)

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Astrophysical black holes have to be electrically neutral:
- $Q = 0$ : Kerr solution (1963)

Other special cases:
- $a = 0$ : Reissner-Nordström solution (1916, 1918)
- $a = 0$ and $Q = 0$ : Schwarzschild solution (1916)
- $a = 0$, $Q = 0$ and $M = 0$ : Minkowski metric (1907)
General definition of a black hole

The textbook definition

[Hawking & Ellis (1973)]

black hole : \[ \mathcal{B} := \mathcal{M} - J^- (\mathcal{I}^+) \]

where

- \((\mathcal{M}, g)\) = asymptotically flat manifold
- \(\mathcal{I}^+\) = future null infinity
- \(J^- (\mathcal{I}^+)\) = causal past of \(\mathcal{I}^+\)

i.e. black hole = region of spacetime from which light rays cannot escape to infinity

event horizon : \( \mathcal{H} := \partial J^- (\mathcal{I}^+) \)

(boundary of \(J^- (\mathcal{I}^+)\))

\(\mathcal{H}\) smooth \(\Rightarrow\) \(\mathcal{H}\) null hypersurface
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event horizon: \( H := \partial J^- (I^+) \)

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In general relativity, a black hole contains a region where the spacetime curvature diverges: the singularity \((NB: this \ is \ not \ the \ primary \ definition \ of \ a \ black \ hole)\). The singularity is inaccessible to observations, being hidden by the event horizon.
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• The singularity marks the limit of validity of general relativity: to describe it, a quantum theory of gravitation would be required.
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The event horizon \(\mathcal{H}\) is a global structure of spacetime: no physical experiment whatsoever can detect the crossing of \(\mathcal{H}\).
Viewed by a distant observer, the horizon approach is perceived with an infinite redshift, or equivalently, by an infinite time dilation.

A black hole is not an infinitely dense object: on the contrary it is made of vacuum (except maybe at the singularity); if one defines its “mean density” by

\[ \bar{\rho} = \frac{M}{(4/3\pi R^3)} , \]

then

- for the Galactic centre BH (Sgr A*) : \( \bar{\rho} \sim 10^6 \text{ kg m}^{-3} \sim 2 \times 10^{-4} \rho_{\text{white dwarf}} \)
- for the BH at the centre of M87 : \( \bar{\rho} \sim 2 \text{ kg m}^{-3} \sim 2 \times 10^{-3} \rho_{\text{water}} \)

\[ \Rightarrow \text{a black hole is a compact object} : \frac{M}{R} \text{ large, not } \frac{M}{R^3} ! \]

Due to the non-linearity of general relativity, black holes can form in spacetimes without any matter, by collapse of gravitational wave packets.
Teleological nature of event horizons

The standard definition of a black hole is highly non-local: determination of $\dot{J}^-(\mathcal{I}^+)$ requires the knowledge of the entire future null infinity. Moreover this is not locally linked with the notion of strong gravitational field:

Example of event horizon in a flat region of spacetime:

Vaidya metric, describing incoming radiation from infinity:

$$ds^2 = - \left(1 - \frac{2m(v)}{r}\right) dv^2 + 2dv dr + r^2 (d\theta^2 + \sin^2 \theta d\varphi^2)$$

with

$m(v) = 0$ for $v < 0$

$dm/dv > 0$ for $0 \leq v \leq v_0$

$m(v) = M_0$ for $v > v_0$

[Ashtekar & Krishnan, LRR 7, 10 (2004)]
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⇒ no local physical experiment can locate the event horizon

[Ashtekar & Krishnan, LRR 7, 10 (2004)]
New paradigm for the theoretical approach to black holes: instead of event horizons, black holes are described by

- trapping horizons (Hayward 1994)
- isolated horizons (Ashtekar et al. 1999)
- dynamical horizons (Ashtekar and Krishnan 2002)
- slowly evolving horizons (Booth and Fairhurst 2004)

All these concepts are local and are based on the notion of trapped surfaces.
The 2000’s: the triumph of numerical relativity

[Caltech/Cornell SXS]
[Scheel et al., PRD 79, 024003 (2009)]
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Three kinds of black holes are known in the Universe:

- **Stellar black holes**: supernova remnants:

  \[ M \sim 10 - 30 \, M_\odot \quad \text{and} \quad R \sim 30 - 90 \, \text{km} \]

  example: Cyg X-1: \( M = 15 \, M_\odot \) and \( R = 45 \, \text{km} \)

- **Supermassive black holes**, in galactic nuclei:

  \[ M \sim 10^5 - 10^{10} \, M_\odot \quad \text{and} \quad R \sim 3 \times 10^5 - 200 \, \text{UA} \]

  example: Sgr A*: \( M = 4.3 \times 10^6 \, M_\odot \) and \( R = 13 \times 10^6 \, \text{km} = 18 \, R_\odot \) = 0.09UA = 1.4 \times \text{radius of Mercury's orbit} \)

- **Intermediate mass black holes**, as ultra-luminous X-ray sources (?):

  \[ M \sim 10^2 - 10^4 \, M_\odot \quad \text{and} \quad R \sim 300 \, \text{km} - 3 \times 10^4 \, \text{km} \]

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Known black holes

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Stellar black holes in X-ray binaries

[McClintock et al. (2011)]
Stellar black holes in X-ray binaries

Dynamically measured masses of black holes in transient low-mass X-ray binaries (right), compared with measured masses of neutron stars (left)

Black holes in the sky

Supermassive black holes in active galactic nuclei (AGN)

Jet emitted by the nucleus of the giant elliptic galaxy M87, at the centre of Virgo cluster [HST]

$M_{\text{BH}} = 3 \times 10^9 M_\odot$

$V_{\text{jet}} \simeq 0.99 c$
The black hole at the centre of our galaxy: Sgr A*

Measure of the mass of Sgr A* black hole by stellar dynamics:

\[ M_{\text{BH}} = 4.3 \times 10^6 M_\odot \]

Orbit of the star S2 around Sgr A*

\[ P = 16 \text{ yr} , \quad r_{\text{per}} = 120 \text{ UA} = 1400 R_\odot , \quad V_{\text{per}} = 0.02 c \]

[Genzel, Eisenhauer & Gillessen, RMP 82, 3121 (2010)]
Can we see a black hole from the Earth?

Angular diameter of the event horizon of a Schwarzschild BH of mass $M$ seen from a distance $d$:

$$\Theta = 6\sqrt{3} \frac{GM}{c^2 d} \simeq 2.60 \frac{2R_S}{d}$$

Image of a thin accretion disk around a Schwarzschild BH

[Vincent, Paumard, Gourgoulhon & Perrin, CQG 28, 225011 (2011)]
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Largest black holes in the Earth’s sky:

- Sgr A* : $\Theta = 53 \mu\text{as}$
- M87 : $\Theta = 21 \mu\text{as}$
- M31 : $\Theta = 20 \mu\text{as}$

Remark: Black holes in X-ray binaries are $\sim 10^5$ times smaller, for $\Theta \propto M/d$
Reaching the $\mu$as resolution with VLBI

Very Large Baseline Interferometry (VLBI) in (sub)millimeter waves

Existing American VLBI network [Doeleman et al. 2011]
Black holes in the sky

Reaching the $\mu$as resolution with VLBI

Very Large Baseline Interferometry (VLBI) in (sub)millimeter waves

The best result so far: VLBI observations at 1.3 mm have shown that the size of the emitting region in Sgr A* is only 37 $\mu$as

[Doeleman et al., Nature 455, 78 (2008)]

Existing American VLBI network [Doeleman et al. 2011]
Black holes in the sky

The near future: the Event Horizon Telescope

To go further:
- shorten the wavelength: \(1.3 \text{ mm} \rightarrow 0.8 \text{ mm}\)
- increase the number of stations; in particular add ALMA

Atacama Large Millimeter Array (ALMA)
part of the Event Horizon Telescope (EHT) to be completed by 2020
August 2015: VLBI observations involving ALMA and VLBA
VLBA and EHT observations of M87

- 40 μas
- EHT beam
- possible BH shadow
- optically-thick region (≥ 21 μas)
- optically-thin region (40 μas)
- jet base of M87 (VLBA at 43GHz)

Near-infrared optical interferometry: GRAVITY

- GRAVITY instrument at VLTI (2016)
- Beam combiner (the four 8 m telescopes + four auxiliary telescopes)
- Astrometric precision on orbits: \(10 \mu\text{as}\)

[Gillessen et al. 2010]
Black holes in the sky

Near-infrared optical interferometry: GRAVITY

July 2015: GRAVITY shipped to Chile and successfully assembled at the Paranal Observatory

Fall 2016: observations have started!

[MPE/GRAVITY team]
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Gravitational waves

Linearization of Einstein equation in weak field: \( g = \eta + h \),
\( \eta = \) Minkowski metric

\[ g = \eta + h, \quad \eta = \text{Minkowski metric} \]

\[ \Rightarrow \text{wave equation:} \quad \square \bar{h} = -\frac{16\pi G}{c^4} T \] (Lorenz gauge)

with \( \square = -\frac{1}{c^2} \frac{\partial^2}{\partial t^2} + \frac{\partial^2}{\partial x^2} + \frac{\partial^2}{\partial y^2} + \frac{\partial^2}{\partial z^2} \), \( \bar{h} = h - \frac{1}{2} h \eta \) and \( h = \text{Trace}(h) \).

1. \( \eta_{\mu\nu} = \text{diag}(-1,1,1,1) \) en Cartesian coordinates
Observing black holes via gravitational waves: a dream come true

Black holes and gravitational waves

Link between black holes and gravitational waves:
Both are spacetime distortions:
- extreme distortions (black holes)
- small distortions (gravitational waves)

In particular, black holes and gravitational waves are both vacuum solutions of Einstein equation
Observing black holes via gravitational waves: a dream come true

Observational evidence for gravitational waves

Emission of gravitational waves by the neutron star binary system PSR B1913+16 (binary pulsar)

→ Observed decay of the orbital period

\[ P = 7 \text{ h 45 min} \] of the binary pulsar PSR B1913+16 produced by the reaction to gravitational radiation → coalescence in 140 millions year.

[Weisber & Taylor (2002)]
Emission of gravitational waves by the neutron star binary system PSR B1913+16 (*binary pulsar*)

\[ P = 7 \, \text{h} \, 45 \, \text{min} \]

observed decay of the orbital period

produced by the *reaction to gravitational radiation*  
lead to coalescence in 140 millions year.

Nobel Prize in Physics to R. Hulse & J. Taylor (1993)
Measurable effects of a gravitational wave passage

Measure of the distance $L$ between two free masses by a “radar” method:

$$L = \frac{1}{2} c(t_2 - t_1)$$

Variation of length $L$ when a gravitational wave passes by:

$$\delta L \sim h L$$

$h = \text{amplitude of the gravitational wave}$

In practice, $h$ is so small that our senses are not sensitive to it:

for the most important astrophysical sources:

$h \sim 10^{-21}$
Observing black holes via gravitational waves: a dream come true

Advanced LIGO detectors

[Abbott et al., PRL 116, 061102 (2016)]
Observing black holes via gravitational waves: a dream come true

Advanced ground-based GW detectors

- Adv. LIGO: started Sept. 2015
- KAGRA (Japan): 2018

Gravitational wave detector VIRGO in Cascina, near Pisa (Italy) [CNRS/INFN]
Observing black holes via gravitational waves: a dream come true

September 14, 2015, 09:50:45 UTC

Hanford, Washington (H1)

Livingston, Louisiana (L1)

Strain ($10^{-21}$)

Numerical relativity
Reconstructed (wavelet)
Reconstructed (template)

Residual

Normalized amplitude

Frequency (Hz)

Time (s)

Éric Gourgoulhon (LUTH)

100 years of black hole physics

UCL, Louvain, 28 Nov. 2016
Observing black holes via gravitational waves: a dream come true

GW150914 event

Signal:
- $\Delta t = 0.2 \text{ s}$
- $f: 35 \rightarrow 250 \text{ Hz}$
- $h_{\text{max}} = 1.0 \times 10^{-21}$

Matched filter:
- $S/N = 24$
- $F_{\text{false}} = 1/203\,000 \text{ yr}$
- $M_1 = 36 \pm 5 \, M_\odot$
- $M_2 = 29 \pm 4 \, M_\odot$
- $d = 410 \pm 180 \, \text{Mpc}$
- $z = 0.09 \pm 0.04$
- $M_{\text{final}} = 62 \pm 4 \, M_\odot$
- $E_{\text{GW}}^{\text{rad}} = 3.0 \pm 0.5 \, M_\odot c^2$
- $a_1 < 0.7$, $a_2 < 0.9$
- $a_{\text{final}} = 0.67 \pm 0.07$

[Abbott et al., PRL 116, 061102 (2016)]
LIGO events in run O1 (12 Sept. 2015 - 19 Jan. 2016)

[Abbott et al., PRX 6, 041015 (2016)]

*NB:* LVT = LIGO-Virgo Trigger (not significant enough to be a detection)
Observing black holes via gravitational waves: a dream come true

LIGO events in run O1 (12 Sept. 2015 - 19 Jan. 2016)

[Abbott et al., PRX 6, 041015 (2016)]
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GW detectors in different bandwidths

[Diagram showing characteristic strain vs. frequency for different gravitational wave detectors and events, including EPTA, SKA, eLISA, Massive binaries, Resolvable galactic binaries, Supermassive binaries, Extreme mass ratio inspirals, Compact binary inspirals, Core collapse supernovae, Pulsars, Virgo, aLIGO (O1), aLIGO.]
Space detector eLISA (ESA)

Interferometric gravitational wave detector in solar orbit

- theme selected by ESA in 2013 for the L3 mission
- launch around 2028
- technology demonstrator LISA Pathfinder launched on 3 December 2015

[eLISA / NGO]
Observing black holes via gravitational waves: a dream come true

eLISA observations of massive binary BH mergers

Signal-to-noise ratio for gravitational waves from the inspiral of a BH binary at $z = 0.5$

Observing black holes via gravitational waves: a dream come true

Detecting gravitational waves by pulsar timing

Le grand radiotélescope de Nançay fête ses 50 ans. © Observatoire de Paris
EPTA results on supermassive BH binaries

EPTA: European Pulsar Timing Array

[Babak et al., MNRAS 455, 1665 (2016)]
Outline

1. A century-old history
2. Black holes in the sky
3. Observing black holes via gravitational waves: a dream come true
4. Testing general relativity with black holes
Is general relativity unique?

Yes if we assume

- a 4-dimensional spacetime
- gravitation only described by a metric tensor $g$
- field equation involving only derivatives of $g$ up to second order
- diffeomorphism invariance
- $\nabla \cdot T = 0$ ($\Rightarrow$ weak equivalence principle)

The above is a consequence of Lovelock theorem (1972).
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However, GR is certainly not the ultimate theory of gravitation:
- it is not a quantum theory
- cosmological constant / dark energy problem

GR is generally considered as a low-energy limit of a more fundamental theory:
- string theory
- loop quantum gravity
- ...
Testing general relativity with black holes

Extensions of general relativity

Higher dimensions

WEP violations

Extra fields

Lovelock theorem

Diff-invar. violations

Nondynamical fields

Dynamical fields (SEP violations)

Massive gravity

Lorentz-violations

Palatini f(R)
Eddington-Born-Infeld

Scalars

Vectors

Tensors

Scalar-tensor, Metric f(R)
Horndeski, galileons
Quadratic gravity, n-DBI

Scalar-tensor, Metric f(R)
Horndeski, galileons
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Einstein-Aether
Horava-Lifshitz

TeVeS
Bimetric gravity

dRGT theory
Massive bimetric gravity

Einstein-Aether
Horava-Lifshitz
n-DBI

[Éric Gourgoulhon (LUTH)]

[Éric Gourgoulhon (LUTH) 100 years of black hole physics UCL, Louvain, 28 Nov. 2016 52 / 57]
Testing general relativity with black holes

Test: are astrophysical black holes Kerr black holes?

- GR $\implies$ Kerr BH (no-hair theorem)
- extension of GR $\implies$ BH may deviate from Kerr
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**Observational tests**

Search for

- stellar orbits deviating from Kerr timelike geodesics (GRAVITY)
- accretion disk spectra different from those arising in Kerr metric (X-ray observatories, e.g. Athena)
- images of the black hole silhouette different from that of a Kerr BH (EHT)
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Need for a good and versatile geodesic integrator to compute timelike geodesics (orbits) and null geodesics (ray-tracing) in any kind of metric
Gyoto code

Main developers: T. Paumard & F. Vincent

- Integration of geodesics in Kerr metric
- Integration of geodesics in any numerically computed 3+1 metric
- Radiative transfer included in optically thin media
- Very modular code (C++)
- Yorick and Python interfaces
- Free software (GPL): http://gyoto.obspm.fr/

[Vincent, Paumard, Gourgoulhon & Perrin, CQG 28, 225011 (2011)]
[Vincent, Gourgoulhon & Novak, CQG 29, 245005 (2012)]
Boson star = localized configurations of a self-gravitating massive complex scalar field $\Phi \equiv \text{“Klein-Gordon geons”}$

[Bonazzola & Pacini (1966), Kaup (1968)]

Boson stars may behave as black-hole mimickers

- Solutions of the Einstein-Klein-Gordon system computed by means of Kadath
  [Grandclément, JCP 229, 3334 (2010)]

- Timelike geodesics computed by means of Gyoto

Zero-angular-momentum orbit around a rotating boson star based on a free scalar field $\Phi = \phi(r, \theta)e^{i(\omega t + 2\varphi)}$
with $\omega = 0.75 \text{ m}/\text{h}$.

[Granclément, Somé & Gourgoulhon, PRD 90, 024068 (2014)]
Kerr BH \( a/M = 0.9 \)

Boson star \( k = 1, \omega = 0.70 \text{m}/\text{h} \)

[Vincent, Meliani, Grandclément, Gourgoulhon & Straub, CQG 33, 105015 (2016)]
Conclusions

After a century marked by the Golden Age (1965-1975), the first astronomical discoveries and the ubiquity of black holes in high-energy astrophysics, black hole physics is very much alive.
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It is entering a new observational era, with the advent of high-angular-resolution telescopes and gravitational wave detectors, which provide unique opportunities to test general relativity in the strong field regime.
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It is entering a new observational era, with the advent of high-angular-resolution telescopes and gravitational wave detectors, which provide unique opportunities to test general relativity in the strong field regime.

The GW150914 event was both the first direct detection of gravitational waves and the first observation of the merger of two black holes — the most dynamical event in relativistic gravity. The waveform was found consistent with general relativity.